

Asymmetric self-interacting dark matter with canonical seesaw

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Abstract

We study the possibility of generating dark matter (DM) and baryon asymmetry of the universe (BAU) simultaneously in an asymmetric DM framework which also alleviates the small-scale structure issues of cold DM (CDM). While thermal relic of such self-interacting DM (SIDM) remains under-abundant due to efficient annihilation into light mediators, a non-zero asymmetry in dark sector can lead to survival of the required DM in the universe. The existence of a light mediator leads to the required self-interactions of DM at small scales while keeping DM properties similar to CDM at large scales. It also ensures that the symmetric DM component annihilates away, leaving the asymmetric part, in the spirit ofogenesis. The particle physics implementation is done in canonical seesaw models of light neutrino mass, connecting it to the origin of DM and BAU. In particular, we consider type-I and type-III seesaw origin of neutrino mass for simplicity and minimality of the field content. We show that the desired self-interactions and relic of DM together with BAU while satisfying relevant constraints lead to strict limits on DM mass $\mathcal{O}(\text{GeV}) \lesssim M_{\text{DM}} \lesssim 460$ GeV. In spite of being high scale seesaw, the models remain verifiable at different experiments including direct, indirect DM search as well as colliders.

I. INTRODUCTION

Dark matter (DM), an enigmatic form of matter devoid of luminous and baryonic compositions, comprises a substantial fraction of the universe, posing a longstanding puzzle in both cosmology and particle physics. Its presence, inferred from gravitational effects [1–4], holds profound implications for our comprehension of the universe's evolution. The prevailing model in cosmology, known as the Λ CDM model, has achieved notable success in elucidating several crucial aspects of the observable universe. Here, Λ represents the cosmological constant or dark energy, while CDM denotes cold DM; a pressureless, collisionless fluid essential for initiating the formation of gravitational potential wells crucial for structure formation. However, persistent discrepancies arise at smaller scales between observations and the predictions of collisionless CDM, referred to as the too-big-to-fail, missing satellite, and core-cusp problems in the literature [5, 6]. Self-interacting DM (SIDM) emerges as a compelling solution to these discrepancies [7–13], with the requisite self-interaction parameterized in terms of the cross-section to mass ratio, typically denoted as $\sigma/m \sim 1 \text{ cm}^2/\text{g}$. DM self-interactions, facilitated by light force carriers, induce not only substantial self-interaction but also engender velocity-dependent interactions, exhibiting enhanced effectiveness in smaller halos characterized by lower velocity dispersion. However, their

efficacy diminishes in larger halos with higher velocity dispersion, consistent with the collisionless picture of CDM [8–11, 14–17]. Nonetheless, the considerable coupling of DM with light mediators, necessary for self-interaction, leads to significant DM annihilation rates, often resulting in a deficit of relic abundance, particularly in the low DM mass regime [18].

On the other hand, observational data from the cosmic microwave background (CMB) and astrophysical observations consistent with the big bang nucleosynthesis (BBN) predictions reveal a remarkable coincidence - the present-day abundances of DM and baryonic matter are very similar, with the ratio of their energy densities being approximately 5 (i.e. $\rho_{\text{DM}} \simeq 5\rho_B$) [4]. More precisely, CMB experiments determine the relic abundance of DM and baryon to be $\Omega_{\text{DM}}h^2 = 0.12 \pm 0.0012$ and $\Omega_B h^2 = 0.02237 \pm 0.00015$ respectively [4]. While there exist different well-motivated frameworks like the weakly interacting massive particle (WIMP) paradigm of DM [19–24] and baryogenesis [25, 26], leptogenesis [27] to explain the BAU, the remarkable similarity $\rho_{\text{DM}} \simeq 5\rho_B$ may be indicative of a different dynamical origin orogenesis mechanism. The asymmetric DM (ADM)¹ paradigm [30–47] is one of the most popular cogenesis mechanisms where there exists an asymmetry in the number density of DM over anti-DM, similar to baryons. Since the asymmetries in dark and visible sectors have a common origin, ADM can naturally explain similar abundances of DM and visible matter. Many of such ADM scenarios rely on the leptogenesis route to BAU and can also explain the origin of light neutrino mass and mixing, another observed phenomena [48] which remain unexplained in

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¹ See [28, 29] for reviews of ADM scenarios.

the standard model (SM).

The ADM paradigm not only provides aogenesis of baryon and DM but can also lead to the correct thermal relic of SIDM with light mediator which otherwise remain thermally under-abundant in the low mass regime [18]. While it is possible to have a hybrid of thermal and non-thermal contribution to SIDM relic [49–52], ADM provides a more minimal setup with other motivations related to cogenesis. A conserved quantum number in the dark sector equivalent to an asymmetry helps in generating the observed DM relic which can not be depleted beyond a certain limit, in spite of efficient annihilation of DM and anti-DM particles into light mediators [53]. It is noteworthy that in ADM scenarios, the observed DM relic density can be obtained without violating the stringent constraints from the CMB and indirect searches on DM annihilation into charged final states or photons [54], in contrast to conventional DM models. This is primarily due to the key feature of ADM whereby it can result in a reduced rate of DM-anti DM annihilation compared to symmetric DM models, owing to the suppressed population of DM antiparticles in the ADM relic density [55, 56].

Drawing inspiration from these considerations, in this article, we explore asymmetric self-interacting DM (ASIDM) which provides cogenesis together with DM self-interactions. While some earlier works [57–59] considered the possibility of ASIDM, we consider a minimal setup within canonical seesaw mechanisms namely, type-I [60–64] and type-III seesaw [65]. While it is possible to implement the idea of ASIDM within type-II seesaw mechanism [64, 66–69] as well, we consider the other two canonical seesaw mechanisms only such that the field content remains minimal. Thus, our setup can accommodate several observed phenomena in the universe namely, non-zero neutrino mass, DM, and the BAU together with providing solutions to the baryon-DM coincidence puzzle and small-scale structure issues of cold DM. We consider the DM relic to arise purely from the asymmetric component while the net lepton asymmetry is converted to the observed baryon asymmetry through electroweak sphaleron processes. We illustrate the generation of the requisite lepton asymmetry and the observed DM relic abundance by considering both weak washout and strong washout scenarios in both the type-I and type-III frameworks. This is achieved by solving the Boltzmann equation, which accounts for all relevant processes involved in asymmetry generation and transfer between the visible and dark sectors. We show that the requirement of DM relic, DM self-interactions and DM-baryon cogenesis together with all relevant phenomenological constraints require DM mass to remain within sub-TeV ballpark $\mathcal{O}(\text{GeV}) \leq M_{\text{DM}} \lesssim 460 \text{ GeV}$. For stable DM, the models can have observable direct detection prospects like DM-nucleon or DM-electron scatterings as well as collider productions of TeV scale particles. While stable ADM does not have promising indirect detection prospects due to inefficient annihilation rates at

late epochs, a long-lived ADM can show up at gamma-ray or neutrino telescopes due to its decay. We find the parameter space of the model consistent with all phenomenological constraints while keeping the detection prospects promising at different experiments.

The structure of the article is as follows. Section II provides a concise overview of DM self-interaction within the most minimal setup. Subsequently, section III outlines the implementation of asymmetric SIDM, detailing its integration into both type-I and type-III seesaw frameworks and exploring the intricacies of cogenesis. In section IV, we discuss the possibility of unstable DM and elucidate the constraints on DM lifetime. Following this, in Section V, we examine the intriguing detection prospects for asymmetric self-interacting DM through both direct and indirect search experiments. Finally, we conclude in section VI with several technical details in the appendices.

II. THERMAL RELIC OF SIDM

We consider DM to be a fermionic singlet, denoted as χ (i.e. $\chi \equiv \text{DM}$), characterized as odd under an additional discrete symmetry Z_2 , which guarantees its stability. One light scalar mediator, represented by Φ , which transforms trivially under Z_2 is invoked to mediate the DM self-interactions. The Lagrangian governing self-interaction can be written as:

$$\mathcal{L}_{\text{self-int.}} = -\lambda_{\text{DM}} \Phi \bar{\chi} \chi + \text{h.c.} \quad (1)$$

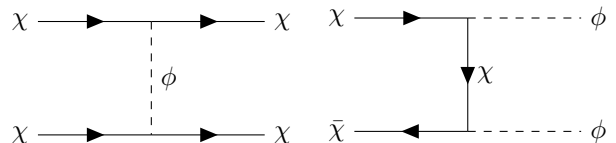


FIG. 1: [left] DM self-interaction mediated by light scalar ϕ and [right] DM annihilation to a pair of light mediator, ϕ , where ϕ is the physical scalar as given in Eq 4.

In this scenario, the non-relativistic self-interaction of DM is effectively described by a Yukawa-type potential $\frac{\lambda_{\text{DM}}^2}{4\pi r} e^{-M_\phi r}$. The details on the non-relativistic self-interaction cross-section is given in Appendix A. The Feynman diagram for DM self-interaction and the dominant number changing process of DM i.e. DM annihilation into the light mediator ϕ , are shown in Fig. 1.

In Fig. 2, we illustrate the parameter space within the $\lambda_{\text{DM}} - M_{\text{DM}}$ plane capable of generating the desired DM self-interactions for two distinct mediator mass values ($M_\phi = 10 \text{ MeV}$ and 100 MeV). The blue dashed line represents the contour of the correct thermal relic abundance of DM, which is solely decided by its annihilation to ϕ , the cross-section for which, ignoring mass of ϕ , is

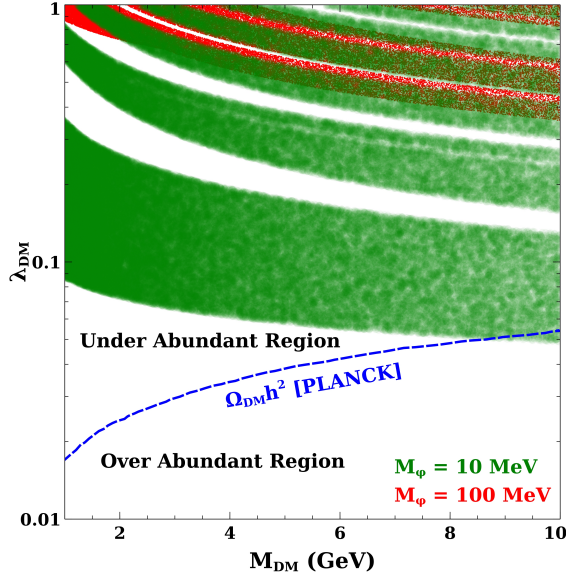


FIG. 2: The parameter space depicted in λ_{DM} vs M_{DM} plane for required self-interaction (colored regions) and correct relic abundance (blue contour) with $M_\phi = 10$ MeV and 100 MeV.

given by

$$\langle \sigma v \rangle_{\chi \bar{\chi} \rightarrow \phi \phi} \sim \frac{\lambda_{\text{DM}}^4}{16\pi M_{\text{DM}}^2}. \quad (2)$$

Evidently, in the low mass regime, the parameter space meeting the self-interaction criteria results in an under-abundant thermal DM relic. This deficit can be avoided in an asymmetric SIDM setup, which we will discuss in the next section.

The scalar potential of the model containing singlet scalar Φ and the SM Higgs doublet H , is given by

$$V(H, \Phi) = \frac{1}{2}\mu_\Phi^2 \Phi^2 - \mu_H^2 (H^\dagger H) + \lambda_H (H^\dagger H)^2 + \frac{\lambda_\Phi}{4} \Phi^4 + \mu_1 \Phi (H^\dagger H) + \frac{1}{2} \lambda_{H\Phi} (H^\dagger H) \Phi^2 \quad (3)$$

Parameterizing the scalars H and Φ as

$$H = \frac{1}{\sqrt{2}} \begin{pmatrix} 0 \\ v_0 + h \end{pmatrix}, \quad \Phi = \phi + u \quad (4)$$

the vacuum expectation values (VEV) are estimated to be

$$v_0 \simeq \sqrt{\frac{2\mu_H^2 - \lambda_{H\Phi} u^2 - 2\mu_1 u}{2\lambda_H}}, \quad u \simeq -\frac{\mu_1 v_0^2}{2\mu_\Phi^2 + \lambda_{H\Phi} v_0^2}.$$

The mass-squared matrix spanning the h and ϕ is given by,

$$M_{h\phi}^2 \simeq \begin{pmatrix} 2\lambda_H v_0^2 & \mu_1 v_0 + \lambda_{H\Phi} u v_0 \\ \mu_1 v_0 + \lambda_{H\Phi} u v_0 & 2\lambda_\Phi u^2 - \frac{1}{2} \frac{\mu_1 v_0^2}{u} \end{pmatrix}, \quad (5)$$

diagonalizing which, we get the mass eigenstates h_1 and h_2 with $h - \phi$ mixing given by

$$\tan 2\gamma \sim \frac{2(\mu_1 v_0 + \lambda_{H\Phi} u v_0)}{2\lambda_H v_0^2 - 2\lambda_\Phi u^2 + \frac{1}{2} \frac{\mu_1 v_0^2}{u}}. \quad (6)$$

After the electroweak symmetry breaking, the SM Higgs mixes with ϕ , thus opening up the window for direct detection of DM.

III. ASYMMETRIC SIDM IN SEESAW FRAMEWORKS

To successfully realize the ASIDM scenario within the type-I seesaw framework, which extends the standard model (SM) by introducing right-handed neutrinos (RHN) N_{R_i} to facilitate light neutrino mass generation, it is imperative to introduce an additional dark singlet scalar (ρ) into the particle spectrum alongside the fermionic DM particle χ . Similarly, within the type-III seesaw scenario, where the SM is extended with $SU(2)_L$ triplet fermions Σ to achieve light neutrino mass generation, ensuring the gauge invariance of the Lagrangian mandates the inclusion of an extra triplet scalar (Δ) to enable the realization of the ASIDM scenario. In both the type-I and type-III frameworks, the simultaneous CP-violating out-of-equilibrium decay of the RHNs (or triplet fermions) into both the visible and dark sectors generates a net lepton asymmetry and an asymmetry in the DM χ . Here we discuss the details of ASIDM implementation in these two canonical seesaw models.

Model-I (Decay of RHN)

The fermion sector of the SM is extended by incorporating three right handed neutrinos ($N_{R_1}, N_{R_2}, N_{R_3}$) in addition to the singlet Dirac fermion (χ) of bare mass M_χ . The right-handed neutrinos are assigned a lepton number of 0 ($L = 0$), while χ is designated a lepton number of 1 ($L = 1$). In order to allow the coupling of χ with RHNs, we introduce a real singlet scalar ρ keeping it heavier than DM. An additional Z_2 symmetry under which χ, ρ are odd while other fields are even, is imposed to ensure the stability of DM. The real singlet scalar Φ , even under Z_2 symmetry, gives rise to DM self-interactions discussed before.

The relevant Lagrangian is given by,

$$\mathcal{L} \supset -\frac{1}{2} M_{N_R} \bar{N}_R^c N_R - \left[y_N \bar{L} \tilde{H} N_R + y_\chi \bar{N}_R \rho \chi + \text{h.c.} \right] \quad (7)$$

The out-of-equilibrium decay of the lightest RHN into LH and $\chi\rho$ can create lepton and DM asymmetries which will be discussed in details below.

Model-II (Decay of fermion triplet)

In this model, we extend the SM with three hyperchargeless fermion triplets ($\Sigma_1, \Sigma_2, \Sigma_3$), one gauge singlet Dirac

fermion χ with bare mass M_χ , and one hyperchargeless triplet scalar Δ . We assign the fermion triplet zero lepton number and the χ lepton number one. To get the velocity-dependent self-interaction of the DM, we include one light scalar mediator, ϕ . An additional discrete symmetry, Z_2 is imposed, under which χ and Δ are odd, while all other fields are even. χ is chosen to be the lightest Z_2 -odd particle ensuring DM stability.

The relevant Lagrangian in this setup is given by,

$$\mathcal{L} \supset -\frac{1}{2}\text{Tr}[M_\Sigma \bar{\Sigma}^c \Sigma] - \left[\sqrt{2}y_\Sigma \bar{L} \Sigma \tilde{H} + y_\chi \text{Tr}[\bar{\Sigma} \Delta \chi] + \text{h.c.} \right] \quad (8)$$

For details of neutrino mass generation in both these scenarios, please refer to Appendix B.

A. ASIDM and Cogenesis

The additional fermion $x(\equiv N_R \text{ or } \Sigma)$ simultaneously decays into both dark sector ($\chi, y(\equiv \rho \text{ or } \Delta)$) and the visible sector (L, H). If the three Sakharov conditions [70] are satisfied, this can lead to the generation of an asymmetry in both the dark sector and the visible sector. The symmetric portion of the DM is depleted through annihilation into light mediator ensured by sizeable DM coupling to mediator Φ required for self-interactions, while the residual asymmetry contributes to the DM relic as asymmetric DM. We consider a mass hierarchy among the fermion mass states as $M_{x_1} < M_{x_2} < M_{x_3}$. Given that x_1 is the lightest among the three, any asymmetry resulting from the decay of x_2 and x_3 at high temperatures will be washed out by the interactions of x_1 efficient at lower temperatures. Hence, for pragmatic purposes, we solely focus on interactions involving x_1 , which will ultimately generate the asymmetry in both sectors. The asymmetry in both sectors originates from the decay of x_1 . However, this produced asymmetry undergoes partial attenuation due to the wash-out processes, such as inverse decay and scatterings that violate lepton number by 1 and 2 units. In addition, one should also incorporate the lepton number conserving scatterings into the analysis to take into account the transfer of asymmetry between the two sectors.

Here it is worth noting that, in the type-III scenario, the presence of additional gauge interactions associated with the fermion triplets solely impacts the number density of these triplets. As these gauge interactions conserve lepton number, they do not directly contribute to

the asymmetry. As outlined in earlier works on type-III seesaw leptogenesis [71, 72], the gauge interactions facilitate the rapid attainment of thermal equilibrium for triplets with masses up to 10^{14} GeV, assuming a dynamical initial abundance. Even if the initial abundance starts at zero, due to these gauge interactions, the system behaves akin to the case with thermal initial abundance. Contrarily, in the type-I scenario, such a phenomenon does not occur. RHNs lack gauge interactions, and thus the asymmetries are influenced by the initial conditions of the RHN number density.

Depending on the Yukawa coupling and M_{Σ_1} , the triplets predominantly maintain equilibrium via either gauge or Yukawa interactions. The Yukawa interactions are parameterized by a parameter known as effective neutrino mass

$$\tilde{m}_1 = \frac{(y_\Sigma^\dagger y_\Sigma)_{11} v_0^2}{2M_{\Sigma_1}}. \quad (9)$$

On the other hand, the rates of gauge interactions depend solely on the electroweak gauge coupling g and M_{Σ_1} . For lower (higher) M_{Σ_1} values, gauge processes maintain equilibrium for a longer (shorter) duration. The prolonged equilibrium of triplets delays their decoupling and subsequent decay, resulting in a lower number density of triplets available for asymmetry production. This poses challenges in realizing the observed baryon asymmetry for very small $\tilde{m}_1$, as in such a scenario, asymmetry suppression due to reduced abundance of the triplet by virtue of gauge interactions becomes prominent. This scenario is referred to as the ‘‘Gauge Regime’’. Conversely, for larger $\tilde{m}_1$ values, equilibrium is established via the inverse decays, leading to its dominance over gauge interactions. In such cases, triplets remain in equilibrium mainly due to Yukawa interactions, hence termed as the ‘‘Yukawa Regime’’. Here, it is worth mentioning that in the gauge regime, the inverse decay washout effect becomes negligible, and similarly, in the Yukawa regime, the gauge annihilation terms can be ignored. Thus, in the Yukawa regime, the efficiency in the type-III setup is similar to that of a decaying RHN in the type-I setup. Additionally, for small values of $\tilde{m}_1$, typically $\tilde{m}_1 < 10^{-3}$, where Yukawa interactions are subdued compared to gauge interactions, the weak washout regime entirely lies within the gauge regime across all masses of M_Σ .

In the type-I seesaw scenario, the Boltzmann equations (BEs) governing the generation of asymmetries as well as the abundance of the N_{R_1} can be written as,

$$\frac{dY_{N_{R_1}}}{dz} = -\frac{\Gamma_D}{H'z}(Y_{N_{R_1}} - Y_{N_{R_1}}^{\text{eq}}) - \frac{(\Gamma_{\Delta L=0} + \Gamma_{\Delta L=1} + \Gamma_{\Delta L=2})}{H'z}(Y_{N_{R_1}} - Y_{N_{R_1}}^{\text{eq}}) \quad (10)$$

$$\begin{aligned} \frac{dY_{\Delta L}}{dz} &= \epsilon_L \frac{\Gamma_D}{H'z} (Y_{N_{R_1}} - Y_{N_{R_1}}^{\text{eq}}) - \left(\frac{1}{2} \frac{\Gamma_D}{H'z} \frac{Y_{N_{R_1}}^{\text{eq}}}{Y_L^{\text{eq}}} \text{Br}_L + \frac{\Gamma_{\Delta L=1}^W + \Gamma_{\Delta L=2}^W}{H'z} \right) Y_{\Delta L} \\ &\quad - \frac{\Gamma_{N_{R_1}}}{H_1} \text{Br}_L \text{Br}_\chi (I_{T_+}(z)(Y_{\Delta L} + Y_{\Delta\chi}) + I_{T_-}(z)(Y_{\Delta L} - Y_{\Delta\chi})) \end{aligned} \quad (11)$$

$$\begin{aligned} \frac{dY_{\Delta\chi}}{dz} &= \epsilon_\chi \frac{\Gamma_D}{H'z} (Y_{N_{R_1}} - Y_{N_{R_1}}^{\text{eq}}) - \left(\frac{1}{2} \frac{\Gamma_D}{H'z} \frac{Y_{N_{R_1}}^{\text{eq}}}{Y_\chi^{\text{eq}}} \text{Br}_\chi + \frac{\Gamma_{\Delta L=1}^W + \Gamma_{\Delta L=2}^W}{H'z} \right) Y_{\Delta\chi} \\ &\quad - \frac{\Gamma_{N_{R_1}}}{H_1} \text{Br}_L \text{Br}_\chi (I_{T_+}(z)(Y_{\Delta\chi} + Y_{\Delta L}) + I_{T_-}(z)(Y_{\Delta\chi} - Y_{\Delta L})) \end{aligned} \quad (12)$$

where, $z = M_{N_{R_1}}/T$, $H' = H_1 z^{-2}$, is the Hubble parameter with $H_1 = 1.66\sqrt{g_*} M_{N_{R_1}}^2 / M_{\text{pl}}$, g_* is the effective relativistic degrees of freedom and M_{pl} is the Planck mass. Here, $Y_p = n_p/n_\gamma$ denotes comoving number density of particle p with n_p, n_γ denoting number density of p and photon number density of the universe respectively. Sim-

ilarly, $Y_{\Delta p} = (n_p - n_{\bar{p}})/n_\gamma$ denotes the comoving density of the asymmetry in p . Br_p denotes the branching ratio of mother particle's decay into particle p .

Similarly, the Boltzmann equations (BEs) governing the asymmetries as well as the abundance of the BSM fermion, Σ_1 are given as,

$$\frac{dY_{\Sigma_1}}{dz} = -\frac{\Gamma_D}{H'z} (Y_{\Sigma_1} - Y_{\Sigma_1}^{\text{eq}}) - \frac{(\Gamma_{\Delta L=0} + \Gamma_{\Delta L=1} + \Gamma_{\Delta L=2})}{H'z} (Y_{N_{R_1}} - Y_{N_{R_1}}^{\text{eq}}) - \frac{\Gamma_A}{H'z} (Y_{\Sigma_1} - Y_{\Sigma_1}^{\text{eq}}) \quad (13)$$

$$\begin{aligned} \frac{dY_{\Delta L}}{dz} &= \epsilon_L \frac{\Gamma_D}{H'z} (Y_{\Sigma_1} - Y_{\Sigma_1}^{\text{eq}}) - \left(\frac{1}{2} \frac{\Gamma_D}{H'z} \frac{Y_{\Sigma_1}^{\text{eq}}}{Y_L^{\text{eq}}} \text{Br}_L + \frac{\Gamma_{\Delta L=1}^W + \Gamma_{\Delta L=2}^W}{H'z} \right) Y_{\Delta L} \\ &\quad - \frac{\Gamma_{\Sigma_1}}{H_1} \text{Br}_L \text{Br}_\chi (I_{T_+}(z)(Y_{\Delta L} + Y_{\Delta\chi}) + I_{T_-}(z)(Y_{\Delta L} - Y_{\Delta\chi})) \end{aligned} \quad (14)$$

$$\begin{aligned} \frac{dY_{\Delta\chi}}{dz} &= \epsilon_\chi \frac{\Gamma_D}{H'z} (Y_{\Sigma_1} - Y_{\Sigma_1}^{\text{eq}}) - \left(\frac{1}{2} \frac{\Gamma_D}{H'z} \frac{Y_{\Sigma_1}^{\text{eq}}}{Y_\chi^{\text{eq}}} \text{Br}_\chi + \frac{\Gamma_{\Delta L=1}^W + \Gamma_{\Delta L=2}^W}{H'z} \right) Y_{\Delta\chi} \\ &\quad - \frac{\Gamma_{\Sigma_1}}{H_1} \text{Br}_L \text{Br}_\chi (I_{T_+}(z)(Y_{\Delta\chi} + Y_{\Delta L}) + I_{T_-}(z)(Y_{\Delta\chi} - Y_{\Delta L})) \end{aligned} \quad (15)$$

where, $z = M_{\Sigma_1}/T$, and definitions of $H', Y_p, Y_{\Delta p}$ remain same as before. In both sets of BEs for type-I and type-III seesaw, $\Gamma_D = \Gamma_x * K_1(z)/K_2(z)$, Γ_x is the total decay width of mother particle x while $K_1(z), K_2(z)$ being the modified Bessel functions of first and second kind respectively. ϵ_L and ϵ_χ respectively, are the visible sector and dark sector CP asymmetry parameters defined as,

$$\epsilon_L = \frac{\Gamma_{x \rightarrow LH} - \Gamma_{x \rightarrow \bar{L}H^\dagger}}{\Gamma_x}; \epsilon_\chi = \frac{\Gamma_{x \rightarrow \chi y} - \Gamma_{x \rightarrow \bar{\chi} y^\dagger}}{\Gamma_x}, \quad (16)$$

In the BEs, $\Gamma_{\Delta L=1}, \Gamma_{\Delta L=2}$ and $\Gamma_{\Delta L=0}$ are the interaction rates for $\Delta L = 1, \Delta L = 2$ and lepton number conserving scatterings respectively that affects the abundance of $N_{R_1}(\Sigma_1)$. The $\Delta L = 0, \Delta L = 1$ and $\Delta L = 2$ processes have been shown explicitly in Appendix C 1, Appendix C 3 and Appendix C 4 respectively.

Γ_A denotes the gauge interactions of the triplet fermion shown in Appendix C 5. Here it is worth noting that $\Gamma_{\Delta L=1}^W, \Gamma_{\Delta L=2}^W$ are the scattering rates of the $\Delta L = 1, 2$ washout processes that contribute to the diminishing of the asymmetries. For instance, if we consider the process, $N_{R_1} N_{R_1} \rightarrow \chi\chi$, then the interaction rate is expressed as $\Gamma = \langle \sigma v \rangle_{N_{R_1} N_{R_1} \rightarrow \chi\chi} \times n_{N_{R_1}}$. This rate governs the evolution of the N_{R_1} abundance. Additionally, since this process violates lepton number by 2 units, it contributes to the washout of the asymmetries. For these washout processes, the interaction rates will depend on the number density of χ . Thus, the interaction rate for the washout process can be written as: $\Gamma^W = \langle \sigma v \rangle_{N_{R_1} N_{R_1} \rightarrow \chi\chi} \times n_\chi^{\text{eq}}$. The last terms in Eq. (11), (12), (14) and (15) represent the $\Delta L = 0$ processes which transfer the asymmetries from one sector to an-

other. These are shown in Appendix C2. We have used calcHEP [73] to calculate the scattering cross-sections. The analytical expressions for I_{T+} and I_{T-} are adopted from [36].

The produced lepton asymmetry then gets converted to the observed baryon asymmetry ($\eta_B \equiv (n_B - n_{\bar{B}})/n_\gamma = (6.1 \pm 0.3) \times 10^{-10}$) by the electroweak sphalerons [74]. In the type-I seesaw case, the baryon asymmetry and the lepton asymmetries are related as:

$$\begin{aligned}\eta_B(z \rightarrow \infty) &= \frac{C_{L \rightarrow B}}{f} Y_{\Delta L}(z \rightarrow \infty) \\ &= \frac{C_{L \rightarrow B}}{f} \epsilon_L \kappa_L(z \rightarrow \infty) Y_{N_{R_1}}^{\text{eq}}(z \rightarrow 0),\end{aligned}\quad (17)$$

Here $C_{L \rightarrow B}$ is the lepton to baryon conversion factor, which in the type-I case is calculated to be -0.518519 and $f = 2387/86$ is the dilution factor calculated assuming standard photon production from the onset of leptogenesis till recombination [75]. For details of the calculation of $C_{L \rightarrow B}$, please refer to Appendix D. Consequently, the required lepton asymmetry falls within the range $\{3.10468 \times 10^{-8}, 3.42586 \times 10^{-8}\}$. Here, κ_L and κ_χ are the visible sector and dark sector efficiency factors defined as,

$$\kappa_L = \frac{Y_{\Delta L}}{\epsilon_L Y_x^{\text{eq}}|_{T \gg M_x}}; \kappa_\chi = \frac{Y_{\Delta \chi}}{\epsilon_\chi Y_x^{\text{eq}}|_{T \gg M_x}} \quad (18)$$

Conversely, in the type-III seesaw scenario, the relation between the baryon asymmetry and the lepton asymmetries is expressed as

$$\begin{aligned}\eta_B(z \rightarrow \infty) &= 3 \frac{C_{L \rightarrow B}}{f} Y_{\Delta L}(z \rightarrow \infty) \\ &= 3 \frac{C_{L \rightarrow B}}{f} \epsilon_L \kappa_L(z \rightarrow \infty) Y_{\Sigma_1}^{\text{eq}}(z \rightarrow 0).\end{aligned}\quad (19)$$

In our model, $C_{L \rightarrow B} = -0.525253$, which is explicitly calculated in Appendix D. Thus, in this case, the required lepton asymmetry ranges from $\{1.02163 \times 10^{-8}, 1.12732 \times 10^{-8}\}$.

Here it is worth noting that, the ratio of the DM relic density and baryon relic density can be expressed as

$$R \equiv \frac{\Omega_{\text{DM}} h^2}{\Omega_B h^2} = \frac{f}{(3)C_{L \rightarrow B}} \frac{M_{\text{DM}}}{m_p} \frac{\epsilon_\chi}{\epsilon_L} \frac{\kappa_\chi}{\kappa_L}, \quad (20)$$

where m_p is the proton mass and the factor 3 in the denominator only appears in the type-III seesaw scenario. As mentioned earlier, given the ratio $R \sim 5$, one might naively anticipate that $M_{\text{DM}} \sim 5m_p$, assuming a similarity in the asymmetry of the number densities between DM and baryons. However, as indicated by Eq. (20), the resulting asymmetries in the visible and dark sectors can vary significantly depending on the CP asymmetry parameters ϵ_χ and ϵ_L and corresponding efficiency parameters. Consequently, there exists the possibility for the DM mass to be viable across a wide range of masses.

We begin by solving the Boltzmann equations for the type-I case in both weak and strong washout regimes. For the weak washout regime, the parameters are varied, and we consider the benchmark point BP1 (Type-I) with the following values: $\{M_{N_{R_1}} = 5 \times 10^{12} \text{ GeV}, y_N = 1.4082 \times 10^{-3}, y_\chi = 3.09584 \times 10^{-3}, \epsilon_L = 7.3 \times 10^{-5}, \epsilon_\chi = 5 \times 10^{-6}, M_\chi = 1 \text{ GeV}\}$. In this case, the branching ratios of N_{R_1} decaying into the dark and visible sectors are 70.7317% and 29.2683%, respectively. We illustrate the evolution of the N_{R_1} abundance and the asymmetries in Fig. 3. Starting with zero initial abundance of the RHN, it is produced through inverse decay and scattering processes. The asymmetries in both sectors begin to build up as soon as a sufficient amount of RHNs is produced. Clearly, due to the chosen Yukawa couplings, the scenario falls into the weak washout regime. We can see that the N_{R_1} abundance reaches equilibrium around $z \sim 6$ and immediately departs from equilibrium. Subsequently, the number density experiences Boltzmann suppression, and the washout effects are no longer effective; hence, the asymmetry saturates.

Similarly, we solve the Boltzmann equations in the strong washout regime and illustrate the evolution of the asymmetries in the visible and dark sectors along with the abundance of RHN in Fig. 4. In this case, the parameters are fixed at $\{M_{N_{R_1}} = 4 \times 10^{11} \text{ GeV}, y_N = 1.144 \times 10^{-2}, y_\chi = 1.26996 \times 10^{-2}, \epsilon_L = 5.8 \times 10^{-7}, \epsilon_\chi = 3.2 \times 10^{-11}, M_\chi = 100 \text{ GeV}\}$. We observe that the RHN is produced through inverse decay and scattering processes and reaches equilibrium around $z \sim 1$. However, due to the larger Yukawa couplings, the RHN remains in equilibrium. In this case, the branching ratios into the visible

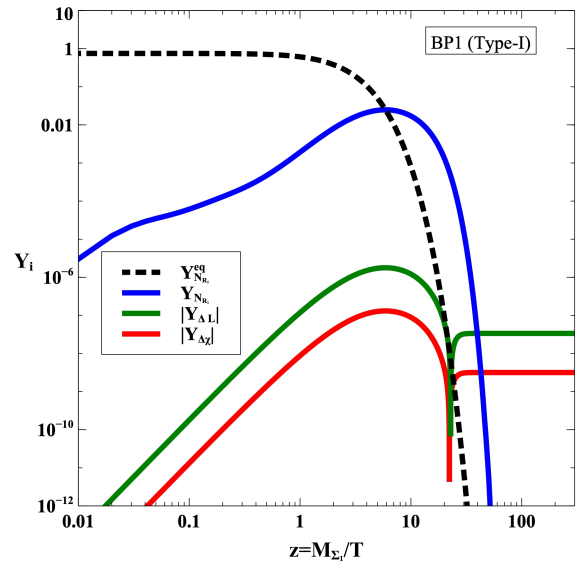


FIG. 3: Cosmological evolution of the RHN abundance, visible sector, and dark sector asymmetries in case of weak washout in type-I seesaw. The final asymmetries are $Y_{\Delta L} = 3.37067 \times 10^{-8}$, $Y_{\Delta \chi} = 3.16549 \times 10^{-9}$, $R = 5.35934$.

and dark sectors are 61.875% and 38.125%, respectively. As the washout processes are strong, the produced asymmetry in the visible sector gets reduced by some factors and finally saturates after $z \sim 20$.

Transitioning to the type-III scenario, we analyze the evolution of the triplet fermion abundance and the resulting asymmetries in both the dark and visible sectors, as depicted in Fig. 5. For this analysis, we set the DM mass at $M_\chi = 1.455$ GeV and establish the parameters as $\{M_{\Sigma_1} = 1.532 \times 10^{12} \text{ GeV}, y_\Sigma = 4.99 \times 10^{-4}, y_\chi = 7.665 \times 10^{-4}, \epsilon_L = 2.9 \times 10^{-6}, \epsilon_\chi = 5.65 \times 10^{-7}\}$, taking into account all relevant processes. As previously mentioned, gauge interactions ensure that the triplets Σ maintain thermal equilibrium. Despite beginning with an initial abundance of zero, the triplets rapidly achieve thermal equilibrium, characterizing this as the gauge regime. In this regime, the triplets remain in equilibrium solely due to gauge interactions. However, around $z \sim 10$, these interactions become subdominant compared to the Hubble rate, leading the triplets to fall out of equilibrium and subsequently decay. By $z \sim 70$, the triplet population is sufficiently depleted, and the asymmetries in both sectors reach their final saturated values. Since this falls within the weak washout regime, the final asymmetries do not experience any suppression.

In Fig. 6, we illustrate the evolution of the triplet fermion abundance and the resulting asymmetries in both sectors for the type-III case within the strong washout regime. The parameters are set as $\{M_{\Sigma_1} = 6.28 \times 10^{12} \text{ GeV}, y_\Sigma = 0.1014, y_\chi = 3.662 \times 10^{-3}, \epsilon_L = 1.4 \times 10^{-6}, \epsilon_\chi = 1.8 \times 10^{-10}, M_\chi = 88.31 \text{ GeV}\}$. This parameter set falls within the Yukawa regime. In this regime, the triplet enters equilibrium rapidly due to

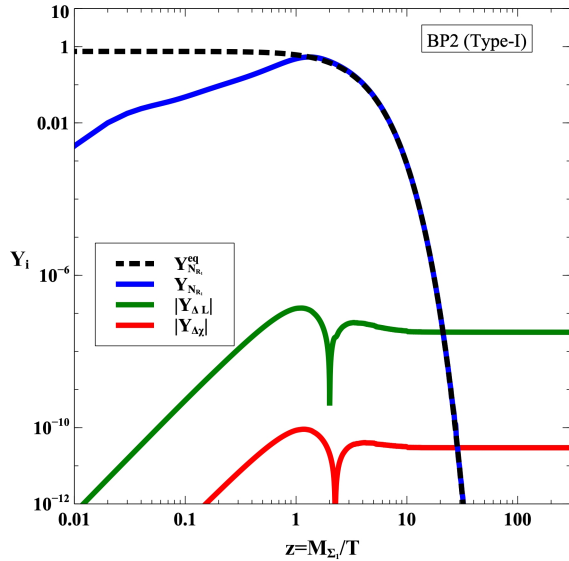


FIG. 4: Cosmological evolution of the RHN abundance, visible sector, and dark sector asymmetries in case of strong washout in type-I seesaw. The final asymmetries are $Y_{\Delta L} = 3.20449 \times 10^{-8}$, $Y_{\Delta \chi} = 2.98107 \times 10^{-11}$, $R = 5.30885$.

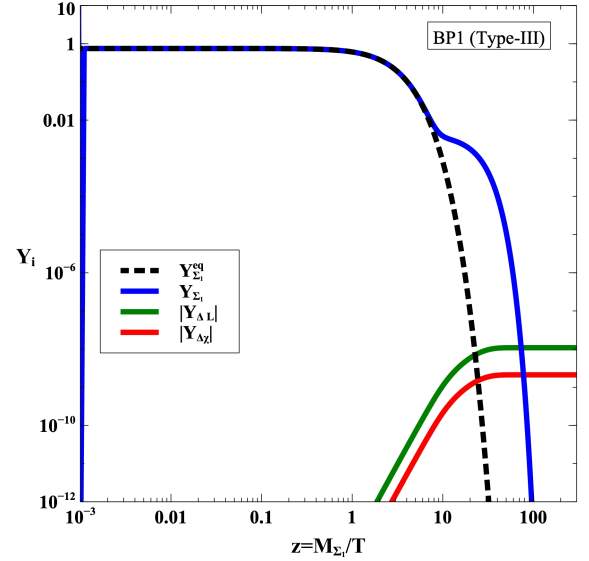


FIG. 5: Cosmological evolution of the fermion triplet Σ_1 abundance, visible sector, and dark sector asymmetries in case of weak washout in type-III seesaw. The final asymmetries are $Y_{\Delta L} = 1.08814 \times 10^{-8}$, $Y_{\Delta \chi} = 2.12 \times 10^{-9}$, $R = 5.32324$.

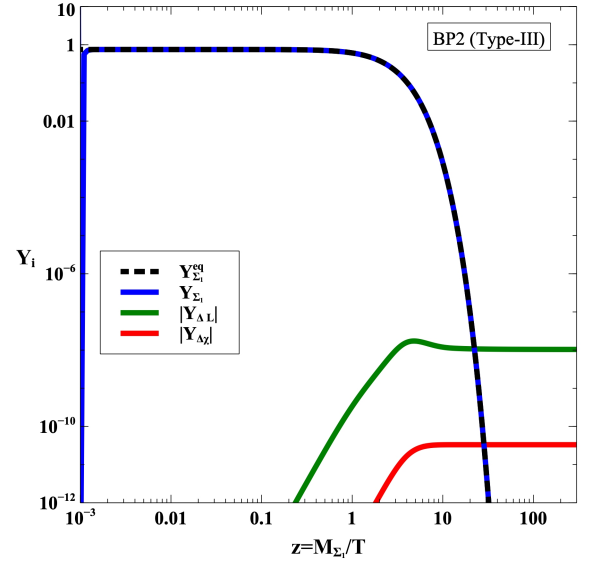


FIG. 6: Cosmological evolution of the triplet abundance, visible sector, and dark sector asymmetries in case of strong washout in type-III seesaw. The final asymmetries are $Y_{\Delta L} = 1.0421 \times 10^{-8}$, $Y_{\Delta \chi} = 3.32243 \times 10^{-11}$, $R = 5.2871$.

gauge interactions, but its subsequent evolution is governed by the Yukawa interactions, which maintain its equilibrium state. In this scenario, the branching ratios for the triplet decaying into the visible and dark sectors are 99.93% and 0.07%, respectively. Due to the substantial branching into the visible sector, the washout effects from inverse decay processes are significant, compounded by additional washouts arising from the lepton

number violating scattering processes. Consequently, we observe a visible suppression in the final lepton asymmetry compared to the dark sector asymmetry, as depicted in Fig. 6.

IV. UNSTABLE ASIDM

In the above discussion, we have imposed an additional discrete Z_2 symmetry to ensure the stability of the DM particle χ , where only χ and the scalar ρ (in the type-I scenario) or Δ (in the type-III scenario) are odd while all other particles transform trivially. However, if we allow soft Z_2 symmetry breaking terms in the scalar potential, the situation changes after electroweak symmetry breaking. When the SM Higgs field acquires a non-zero VEV, the scalars ρ or Δ also acquire an induced VEV due to the presence of interaction terms such as $\mu_\rho \rho H^\dagger H$ or $\mu_\Delta H^\dagger (\vec{\sigma} \vec{\Delta} H)$, which softly break the Z_2 symmetry. This not only opens up new detection prospects of DM, but can also erase the asymmetries associated with the scalars, such that they do not affect visible and DM asymmetries via late decay. This ensures DM relic to be decided by the asymmetry generated at seesaw scale with the symmetric part getting annihilated away. Here, it is worth noting that the annihilation of symmetric component of DM freezes out well before BBN.

As mentioned above, the soft breaking of the Z_2 symmetry has crucial consequences regarding the stability of the DM particle χ , opening up its decay modes. When the scalars ρ (Δ) in Model-I (Model-II) acquire non-zero vacuum expectation values, it induces a small mixing between χ and the right-handed neutrino N_{R_1} (or the triplet fermion Σ_1). Consequently, due to the seesaw mechanism, this mixing propagates to the light neutrino sector. As a result, an effective mixing between the DM particle χ and the light neutrinos is generated. This effective mixing between χ and neutrinos opens up various decay channels for the DM particle, including νZ , $e^- W^+$, $\phi \nu$, and $\nu \bar{f} f$. The Feynman diagrams illustrating these decay modes are shown in Fig. 7, the decay rates for which are given in Appendix E.

For χ to qualify as a viable DM candidate, its lifetime (τ_{DM}) should be greater than 10^{27} seconds [76], which is a conservative lower limit obtained using the gamma-ray data from the Fermi-LAT observation of Milky Way satellite dSphs. The dominant decay modes of the DM particle depend on its mass range. When the DM mass lies between 100 MeV and M_W , it predominantly decays into neutrinos (ν) and scalar particles (ϕ). In this range, the decay width is inversely proportional to the DM mass. Consequently, as the DM mass increases, its lifetime also increases. Thus as the DM lifetime is inversely proportional to the square of the VEV v_1 , to maintain a constant lifetime, the value of v_1 must increase as the DM mass increases. This is evident from Fig. 8, where we have shown the contours of different values of DM lifetime in the plane of DM mass and VEV v_1 . In the mass range be-

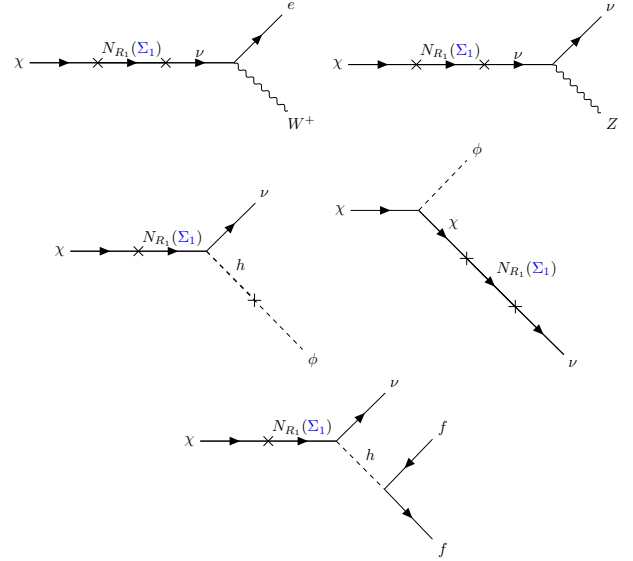


FIG. 7: Feynman diagram of different decay modes of χ .

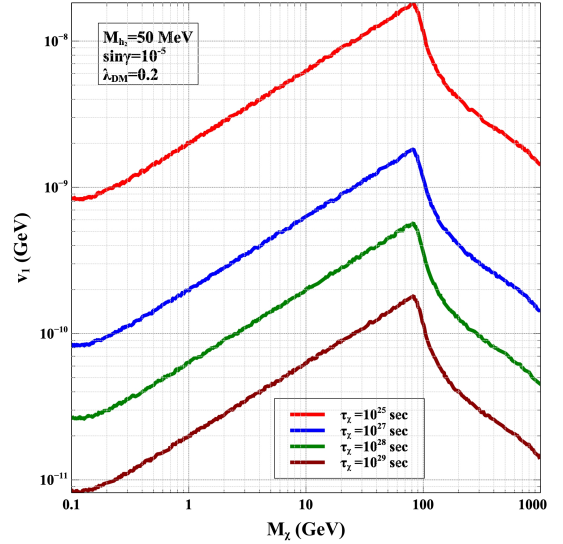


FIG. 8: VEV of the scalar, ρ (Δ) versus DM mass for fixed DM lifetime. Here we have kept $y_\chi = 10^{-3}$, $M_{R_1}(M_{\Sigma_1}) = 10^{12}$ GeV, $M_{h_2} = 50$ MeV, $\lambda_{\text{DM}} = 0.2$, and $\sin \gamma = 1 \times 10^{-5}$.

tween M_W and M_Z , the DM predominantly decays into W^+ and e^- . In this case, the decay width is proportional to the DM mass. As the DM mass increases, the lifetime increases, and to maintain a fixed lifetime, the value of v_1 must decrease. When the DM mass falls within the range of M_Z to M_{h_1} , the dominant decay mode is into Z and ν . Here, the decay width is also proportional to the DM mass, and a similar behavior is observed as in the previous range. Once the DM mass exceeds the Higgs mass, the h and ν decay channel opens up, further increasing the decay width. Consequently, the value of v_1 decreases as the DM mass increases in this regime. These behaviors are prominently illustrated in Fig. 8. Clearly,

the parameter space below the blue contour remains safe from this contour. Thus depending on the value of the v_1 , a wide range of DM masses can be accommodated while satisfying the DM lifetime constraint.

V. DETECTION PROSPECTS

In this section, we discuss the prospects of detecting ASIDM of our frameworks in direct and indirect DM search experiments.

A. Direct Detection

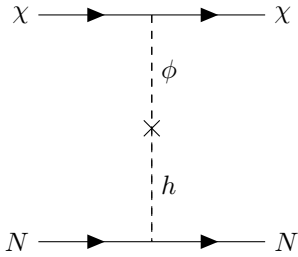


FIG. 9: The spin-independent scattering of DM-nucleon (N) via Higgs portal.

The possibility of spin-independent (SI) DM-nucleon elastic scattering allows for the detection of DM in terrestrial laboratories. In this setup, SI elastic scattering of DM is achievable in terrestrial laboratories via $h - \phi$ mixing, where DM particles can scatter off target nuclei. The Feynman diagram for this process is shown in Fig. 9.

The spin-independent elastic scattering cross-section of DM per nucleon can be calculated as [77]

$$\sigma_{\text{DM-N}}^{\text{SI}} = \frac{\mu_r^2}{\pi A^2} [Z f_p + (A - Z) f_n]^2 \quad (21)$$

where $\mu_r = \frac{M_\chi m_n}{M_\chi + m_n}$ is the reduced mass, m_n is the nucleon (proton or neutron) mass, A is the mass number of target nucleus, Z is the atomic number of target nucleus. The f_p and f_n are the interaction strengths of proton and neutron with DM, respectively, and are given as,

$$f_{p,n} = \sum_{q=u,d,s} f_{T_q}^{p,n} \alpha_q \frac{m_{p,n}}{m_q} + \frac{2}{27} f_{T_G}^{p,n} \sum_{q=c,t,b} \alpha_q \frac{m_{p,n}}{m_q} \quad (22)$$

where,

$$\alpha_q = \lambda_{\text{DM}} * \frac{m_q}{v_0} * \sin(\gamma) \cos(\gamma) \left(\frac{1}{M_{h_2}^2} - \frac{1}{M_{h_1}^2} \right) \quad (23)$$

with γ being the mixing angle between h and ϕ . The values of $f_{T_q}^{p,n}$, $f_{T_G}^{p,n}$ can be found in [78].

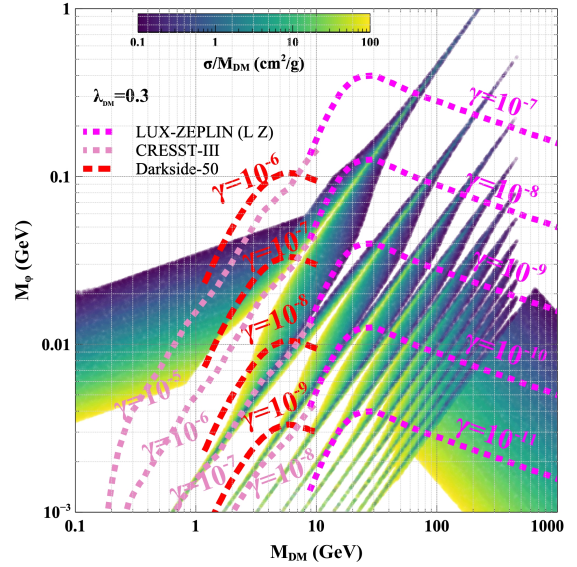


FIG. 10: DM self-interaction preferred region in the plane of $M_{\text{DM}} - M_\phi$ for a fixed value of $\lambda_{\text{DM}} = 0.3$. Direct detection constraints from the LZ, CRESST-III experiments along with future sensitivity of DarkSide-50 are shown as different contours.

Direct search experiments like the CRESST-III [79] and LUX-ZEPLIN (LZ) [80], put severe constraints on the model parameters. The LZ experiment provides the most stringent constraint on DM mass above 10 GeV, while the CRESST-III data constrain the mass regime below 10 GeV. In Fig. 10, the most stringent constraints from the CRESST-III [79], LZ [80] and the anticipated sensitivity of DarkSide-50 [81] experiments are shown against the self-interaction favored parameter space in the plane of M_{DM} and M_ϕ assuming a fixed value of $\lambda_{\text{DM}} = 0.3$. The region below each of these contours is excluded for that particular value of the mixing angle γ .

The mixing angle γ is subject to constraints from both upper and lower bounds. The upper bound arises from the consideration of invisible Higgs decay, as the singlet scalar is typically lighter than the Higgs mass. On the other hand, a lower bound on γ can be obtained by requiring the scalar particle ϕ to decay before the epoch of BBN, i.e., $\tau_\phi < \tau_{\text{BBN}}$, where τ_ϕ and τ_{BBN} represent the lifetime of ϕ and the BBN timescale, respectively. This condition ensures that the decay of ϕ does not disrupt the successful predictions of primordial nucleosynthesis.

In Fig. 11, we have shown the allowed parameter space from the direct detection, DM self-interaction, and BBN constraints in the plane of $\sin \gamma$ and M_ϕ for a fixed DM mass, $M_{\text{DM}} = 5$ GeV, and coupling, $\lambda_{\text{DM}} = 1$. The blue-shaded region is excluded by the CRESST-III [79] experiment due to the constraints on the DM-nucleon spin-independent scattering cross-section. The grey-shaded region is disallowed by the BBN constraint on the ϕ decay lifetime, i.e., in this region, $\tau_\phi > \tau_{\text{BBN}}$. The green region is excluded due to the requirement of sufficient self-interaction among the DM.

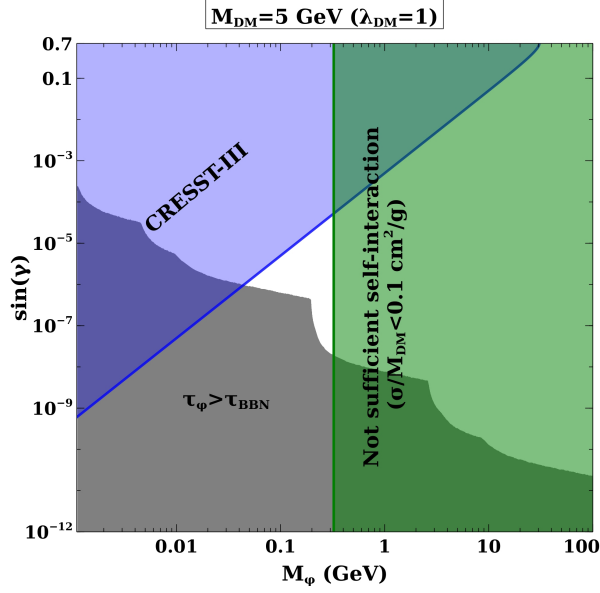


FIG. 11: Constraints from DM direct detection, DM self-interaction, and BBN in the plane of $\sin(\gamma) - M_\phi$ for DM mass of 5 GeV.

To capture the final allowed parameter space, we have performed a scan to simultaneously satisfy the constraints from DM self-interaction, direct detection, and BBN constraints on light mediator mass and mixing angle γ . The result of this scan is shown in Figs. 12 and 13. We have varied the free parameters within the following ranges: $M_{\text{DM}} \in [0.1, 10^3]$ GeV, $M_\phi \in [10^{-3}, 10]$ GeV, $\sin \gamma \in [10^{-12}, 0.7]$, and $\lambda_{\text{DM}} \in [0.1, \sqrt{4\pi}]$. We show the parameter space in the plane of M_{DM} and M_ϕ in Fig. 12, with the color code representing λ_{DM} . The results depict that the maximum allowed DM mass that can satisfy all the constraints is around $M_{\text{DM}} \sim 460$ GeV. In Fig. 13, we represent the parameter space in the plane of $\sin \gamma$ and M_ϕ , with the DM mass shown in the color code. Clearly, a smaller mixing angle, γ , for small M_ϕ is disfavored by the BBN constraint, as it would lead to a lifetime of ϕ greater than τ_{BBN} . We also observe that if DM mass is light then large $\sin \gamma$ is allowed because of weaker constraints from DD constraints which can be probed by the future experiments with enhanced sensitivities. The upper limit on M_{DM} can be understood from the following reasoning. As evident from Fig. 10, for DM masses greater than 460 GeV, the correct self-interaction cross-section can only be obtained if the mediator mass (M_ϕ) is very light, below approximately 20 MeV, falling within the classical regime as elaborated in Appendix A. However, such a small mediator mass results in a sizable DM-nucleon scattering cross-section which is ruled out by the direct detection constraints. Consequently, to adhere to the constraints from direct detection, the mixing angle γ must be decreased. Nevertheless, these small values of M_ϕ and γ are excluded by the BBN constraints on the lifetime of ϕ . Therefore, the simultaneous imposition

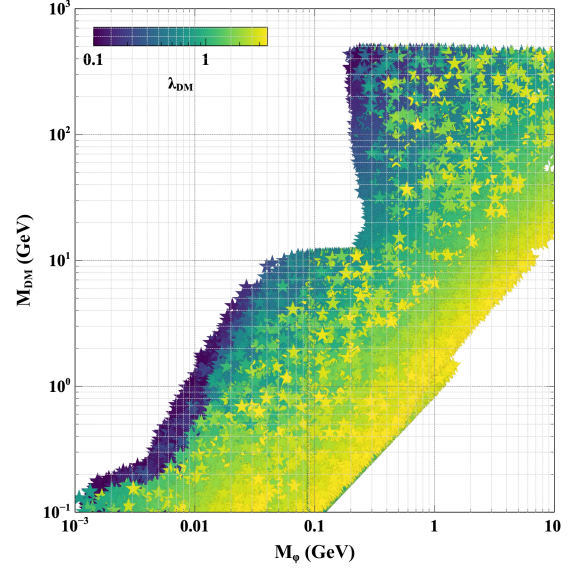


FIG. 12: Parameter space in the plane of M_{DM} and M_ϕ that can give rise to the required self-interaction among the DM while being consistent with the direct detection and BBN constraints.

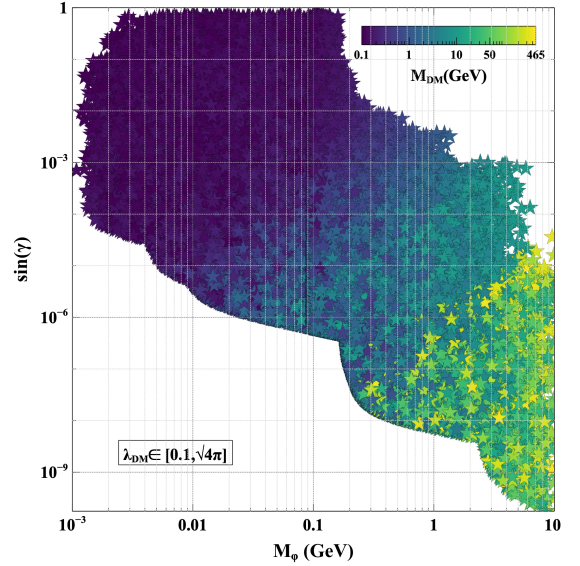


FIG. 13: Parameter space in the plane of M_ϕ and $\sin \gamma$ which is consistent with the direct detection and BBN constraints while still giving rise to the required self-interaction.

of constraints from DM self-interaction, direct detection experiments, and the BBN constraint on the lifetime of the light mediator limits the DM mass to be less than 460 GeV.

B. Indirect detection

As previously discussed, due to the effective mixing between the DM and active neutrinos, a distinct decay mode for the DM can occur, resulting in the emission of a monochromatic gamma-ray line at energy $E = M_{DM}/2$. This emission arises from a one-loop decay process, $\chi \rightarrow \nu\gamma$, as depicted in Fig. 14. The corresponding decay width can be estimated as [82, 83]

$$\tau_{\chi \rightarrow \nu\gamma} \simeq \left(\frac{9G_F^2 \alpha_{EM}}{1024\pi^4} \sin(\theta_{\chi x})^2 \sin(\theta_{\nu x})^2 M_\chi^5 \right)^{-1}, \quad (24)$$

where $\alpha_{EM} = 1/137$ is the fine structure constant,

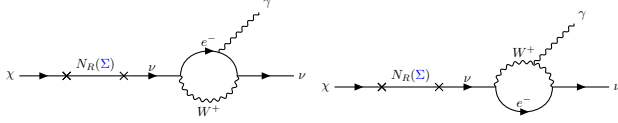


FIG. 14: One-loop decay of DM into $\nu\gamma$.

$G_F = 1.166 \times 10^{-5} \text{ GeV}^{-2}$ is the Fermi constant, $\theta_{\chi x}$ is the mixing angle between χ and $N_R(\Sigma)$, and $\theta_{\nu x}$ is the mixing angle between ν and $N_R(\Sigma)$. In Fig. 15, we show the constraints from the gamma-ray search by the Fermi-LAT [84] and DAMPE [85] experiments in the $\sin(\theta_{\chi x}) - \sin(\theta_{\nu x})$ plane for different DM masses. The parameter space below these contours remain safe from such gamma-ray constraints.

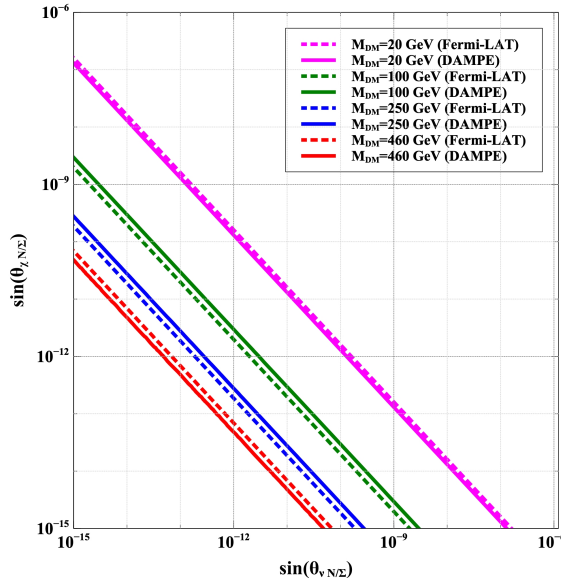


FIG. 15: Constraints from the Fermi-LAT and DAMPE experiments on monochromatic gamma-ray emissions from DM in the plane of $\sin(\theta_{\chi x}) - \sin(\theta_{\nu x})$ for different DM masses. The parameter space towards the bottom left part of each contour remains allowed for respective DM mass.

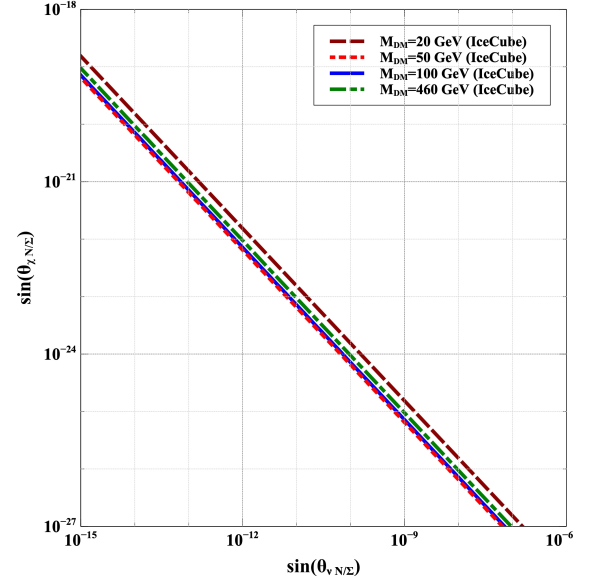


FIG. 16: Constraints from the IceCube experiment on neutrino lines from DM in the plane of $\sin(\theta_{\chi x}) - \sin(\theta_{\nu x})$ for different DM masses. The parameter space towards the bottom left part of each contour remains allowed for respective DM mass.

Moreover, two-body decay of the DM into neutrinos like $\chi \rightarrow \nu\phi$ in our model can give rise to monochromatic neutrino lines. This decay width is given by

$$\tau_{\chi \rightarrow \nu\phi} = \left(\frac{1}{16\pi} \lambda_{DM}^2 \sin(\theta_{\chi x})^2 \sin(\theta_{\nu x})^2 M_{N(\Sigma)}^2 \frac{1}{M_\chi} \right)^{-1} \quad (25)$$

In Fig. 16, we show the constraints from the neutrino line searches by the IceCube [86] experiment in the $\sin(\theta_{\chi x}) - \sin(\theta_{\nu x})$ plane. We show contours for four different DM masses $\{20 \text{ GeV}, 50 \text{ GeV}, 100 \text{ GeV}, 460 \text{ GeV}\}$. We see that when the mixing angle $\sin(\theta_{\nu x}) \sim 10^{-7}$, the $\chi - x$ mixing angle has to be $\sin(\theta_{\chi x}) < 10^{-27}$. In Fig. 15, we see that when the mixing angle $\sin(\theta_{\nu x}) \sim 10^{-7}$, the $\chi - x$ mixing angle has to be $\sin(\theta_{\chi x}) < 10^{-15}$ for $M_{DM} = 20 \text{ GeV}$. Thus the neutrino lines from the IceCube gives more stringent constraint on the mixing angle than the ones from gamma-ray searches by the Fermi-LAT or DAMPE experiment.

VI. CONCLUSION

In this study, we investigate DM self-interaction and the baryon-DM coincidence problem within the context of two popular canonical seesaw frameworks, namely type-I and type-III seesaw, which explain non-zero neutrino masses. While DM self-interactions facilitated by a light mediator offer a potential solution to small-scale astrophysical structure issues of CDM, they often lead to

thermally under-abundant DM relic due to strong annihilation rates of DM into light mediators. To address this challenge as well as to explain the ratio of DM to baryon density, we pursue the idea of asymmetric DM andogenesis within type-I and type-III seesaw models. The CP-violating decays of RHNs (or fermion triplets) into SM leptons and the Higgs boson can produce a lepton asymmetry, while their decay into DM and a singlet scalar ρ (or triplet scalar Δ) simultaneously generates a DM asymmetry. This lepton asymmetry is then transformed into a baryon asymmetry, with the DM asymmetry persisting as the DM relic. DM self-interactions also ensure that the symmetric DM component annihilates away leaving only the asymmetric part, similar to baryons. While the models have promising detection prospects at direct search experiments as well as astrophysical surveys, the conventional indirect detection prospects remain low due to absence of efficient annihilation rates of asymmetric DM in the present universe. This can however change if the Z_2 symmetry protecting DM stability is broken softly, leading to induced VEVs of Z_2 -odd scalars thereby opening up DM decay channels into SM particles. By adjusting the induced VEVs of these scalars, it is possible to ensure the longevity of DM, making it compatible with existing DM lifetime constraints. The decay rate of DM can already saturate bounds on mono-chromatic gamma rays and neutrinos keeping the indirect detection prospects promising. Through an extensive analysis of the model parameter space, we identify the essential conditions for achieving successful cogenesis and adequate self-interaction while adhering to stringent phenomenological and experimental constraints. Our results demonstrate that the model is highly predictive and amenable to testing through a variety of direct and indirect DM detection.

Acknowledgments

S.M. acknowledges the financial support from National Research Foundation(NRF) grant funded by the Korea government (MEST) NRF-2022R1A2C1005050. The work of D.B. is supported by the Science and Engineering Research Board (SERB), Government of India grants MTR/2022/000575 and CRG/2022/000603. P.K.P. would like to acknowledge the Ministry of Education, Government of India, for providing financial support for his research via the Prime Minister's Research Fellowship (PMRF) scheme. The works of N.S. and P.S. are supported by the Department of Atomic Energy-Board of Research in Nuclear Sciences, Government of India (Ref. Number: 58/14/15/2021- BRNS/37220).

Appendix A: DM self-interaction cross-section

The self-interaction of non-relativistic DM can be described by a Yukawa potential given as,

$$V(r) = -\frac{\alpha_D}{r} e^{-M_\phi r} \quad (A1)$$

where $\alpha_D = \lambda_{\text{DM}}^2/4\pi$, and M_ϕ is the mass of the singlet scalar, ϕ . To capture the relevant physics of forward scattering divergence for the self-interaction, we define the transfer cross-section σ_T as [5, 9, 17]

$$\sigma_T = \int d\Omega (1 - \cos\theta) \frac{d\sigma}{d\Omega}. \quad (A2)$$

In the Born Limit ($\alpha_D M_\chi/M_\phi \ll 1$), for both attractive as well as repulsive potentials, the transfer cross-section is

$$\sigma_T^{\text{Born}} = \frac{8\pi\alpha_D^2}{M_\chi^2 v^4} \left(\ln(1 + M_\chi^2 v^2/M_\phi^2) - \frac{M_\chi^2 v^2}{M_\phi^2 + M_\chi^2 v^2} \right). \quad (A3)$$

Outside the Born regime ($\alpha_D M_\chi/M_\phi \geq 1$), we have two distinct regions *viz* the classical region and the resonance region. In the classical limit ($M_\chi v/M_\phi \geq 1$), the solutions for an attractive potential is given by [17, 87, 88]

$$\sigma_T^{\text{classical}} = \begin{cases} \frac{4\pi}{M_\phi^2} \beta^2 \ln(1 + \beta^{-1}) & , \beta \leq 10^{-1} \\ \frac{8\pi}{M_\phi^2} \beta^2 / (1 + 1.5\beta^{1.65}), 10^{-1} \leq \beta \leq 10^3 \\ \frac{\pi}{M_\phi^2} (\ln \beta + 1 - \frac{1}{2} \ln^{-1} \beta)^2, \beta \geq 10^3 \end{cases} \quad (A4)$$

where $\beta = 2\alpha_D M_\phi/(M_\chi v^2)$.

In the resonance region ($M_\chi v/M_\phi \leq 1$), the Yukawa potential is described by a Hulthen potential ($V(r) = -\frac{\alpha_D \delta e^{-\delta r}}{1 - e^{-\delta r}}$) and the solution is given as [17],

$$\sigma_T^{\text{Hulthen}} = \frac{16\pi \sin^2 \delta_0}{M_\chi^2 v^2} \quad (A5)$$

where δ_0 is given by,

$$\delta_0 = \arg \left(\frac{i\Gamma\left(\frac{iM_\chi v}{kM_\phi}\right)}{\Gamma(\lambda_+)\Gamma(\lambda_-)} \right), \lambda_\pm = 1 + \frac{iM_\chi v}{2kM_\phi} \pm \sqrt{\frac{\alpha_D M_\chi}{kM_\phi} - \frac{M_\chi^2 v^2}{4k^2 M_\phi^2}} \quad (A6)$$

with $k \approx 1.6$ being a dimensionless number.

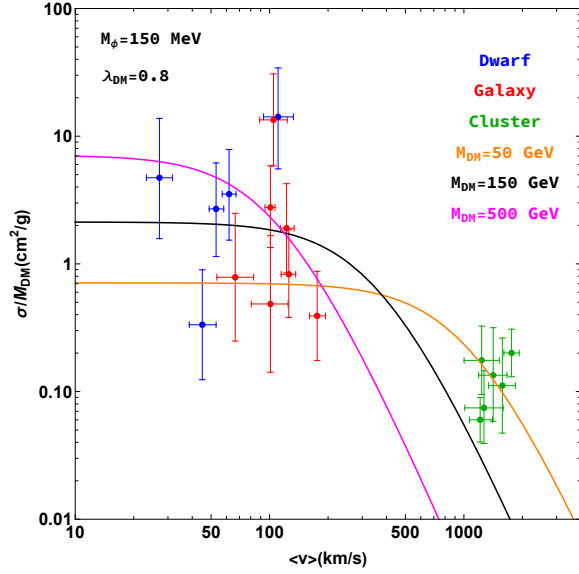


FIG. 17: Self-interacting cross-section per unit DM mass as a function of averaged collision velocity.

We have shown the self-interacting cross-section per unit DM mass as a function of averaged collision velocity for three different DM masses with a fixed coupling of $\lambda_{\text{DM}}=0.8$ and mediator mass, $M_\phi=150$ MeV in Fig. 17 which is consistent with the observational data [15].

Appendix B: Neutrino mass generation

1. Model-I

The Lagrangian responsible for neutrino mass generation is

$$-\mathcal{L}_{\nu\text{-mass}} \supset \frac{1}{2} M_{N_R} \bar{N}_R^c N_R + y_N \bar{L} \tilde{H} N_R + \text{h.c.} \quad (\text{B1})$$

In the effective theory, the masses of light neutrinos can be obtained as:

$$-\mathcal{L}_{\nu\text{-mass}} = (\bar{\nu}_L \quad \bar{N}_R^c) \begin{pmatrix} 0 & \frac{y_N v_0}{\sqrt{2}} \\ \frac{y_N v_0}{\sqrt{2}} & M_{N_R} \end{pmatrix} \begin{pmatrix} \nu_L^c \\ N_R^c \end{pmatrix}. \quad (\text{B2})$$

Diagonalising the above mass matrix we get the Majorana mass of light neutrino as:

$$m_\nu \simeq -\frac{y_N^2 v_0^2}{2M_{N_R}}. \quad (\text{B3})$$

Considering 3 generations of the heavy RHN, the light neutrino mass matrix is given by

$$(m_\nu)_{\alpha\beta} = (m_D)_{\alpha i} (M^{-1})_{ij} (m_D^T)_{j\beta}, \quad (\text{B4})$$

where $m_D = \frac{y_N v_0}{\sqrt{2}}$ is the 3×3 Dirac mass matrix and M is the 3×3 RHN mass matrix. Using the Casas-Ibarra parametrization [89], we can calculate the Yukawa coupling matrix as

$$y_N = \frac{\sqrt{2}}{v_0} (U_{\text{PMNS}} \cdot \sqrt{\hat{m}_\nu} \cdot \mathcal{R} \cdot \sqrt{\hat{M}_N}), \quad (\text{B5})$$

where U_{PMNS} is the lepton mixing matrix, $\hat{m}_\nu$ is 3×3 diagonal light neutrino mass matrix with eigen values m_1, m_2 , and m_3 ; $\hat{M}_N$ is 3×3 diagonal RHN mass matrix with eigen values $M_{N_{R_1}}, M_{N_{R_2}}$, and $M_{N_{R_3}}$, and $\mathcal{R}$ is an arbitrary complex orthogonal matrix.

2. Model-II

The Lagrangian responsible for neutrino mass generation is

$$-\mathcal{L}_{\nu\text{-mass}} \supset \frac{1}{2} M_\Sigma \text{Tr}[\bar{\Sigma}^c \Sigma] + \sqrt{2} y_\Sigma \bar{L} \tilde{H} \Sigma + \text{h.c.} \quad (\text{B6})$$

In the effective theory, the masses of light neutrinos can be obtained as:

$$-\mathcal{L}_{\nu\text{-mass}} = (\bar{\nu}_L \quad \bar{\Sigma}^c) \begin{pmatrix} 0 & \frac{y_\Sigma v_0}{\sqrt{2}} \\ \frac{y_\Sigma v_0}{\sqrt{2}} & M_\Sigma \end{pmatrix} \begin{pmatrix} (\nu_L)^c \\ \Sigma^c \end{pmatrix}. \quad (\text{B7})$$

Diagonalising the above mass matrix we get the Majorana mass of light neutrino as:

$$m_\nu \simeq -\frac{y_\Sigma^2 v_0^2}{2M_\Sigma}. \quad (\text{B8})$$

Considering 3 generations of the heavy fermion triplets, the light neutrino mass matrix is given by

$$(m_\nu)_{\alpha\beta} = (m_D)_{\alpha i} (M^{-1})_{ij} (m_D^T)_{j\beta}, \quad (\text{B9})$$

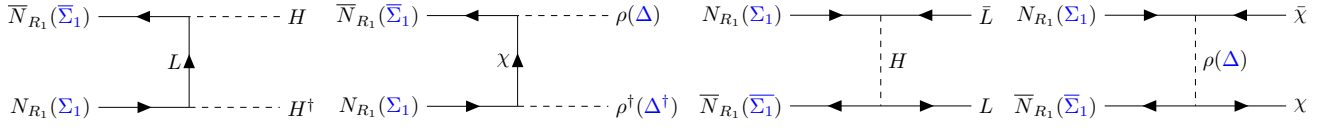
where $m_D = \frac{y_\Sigma v_0}{\sqrt{2}}$ is the 3×3 Dirac mass matrix and M is the 3×3 triplet fermion mass matrix. Using the Casas-Ibarra parameterization [89] we can calculate the Yukawa coupling matrix as

$$y_\Sigma = \frac{\sqrt{2}}{v_0} (U_{\text{PMNS}} \cdot \sqrt{\hat{m}_\nu} \cdot \mathcal{R} \cdot \sqrt{\hat{M}_\Sigma}), \quad (\text{B10})$$

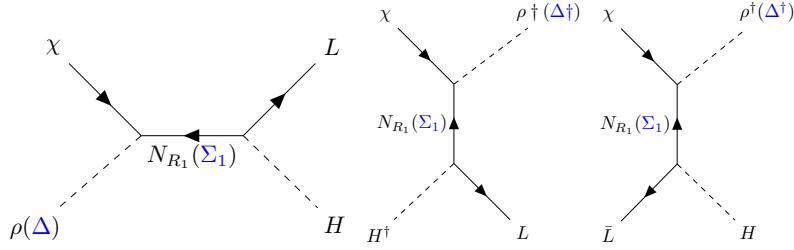
where $\hat{M}_\Sigma$ is 3×3 diagonal triplet fermion mass matrix with eigen values $M_{\Sigma_1}, M_{\Sigma_2}$, and M_{Σ_3} .

Appendix C: Feynmann diagrams for leptogenesis

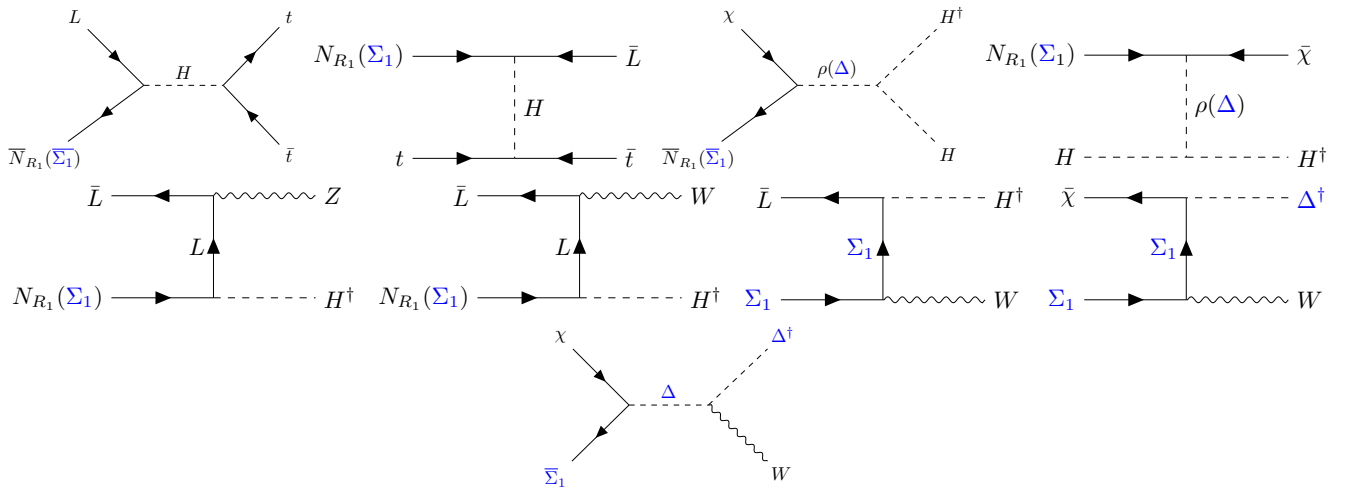
1. $\Delta L = 0$



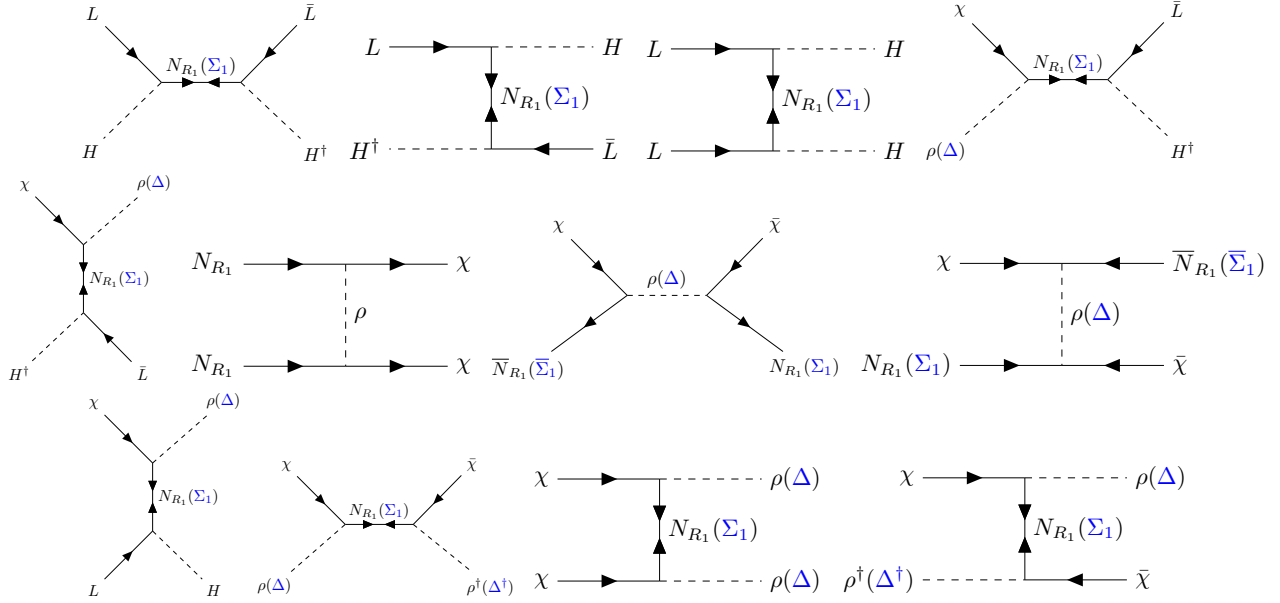
2. $\Delta L = 0$, transfer processes



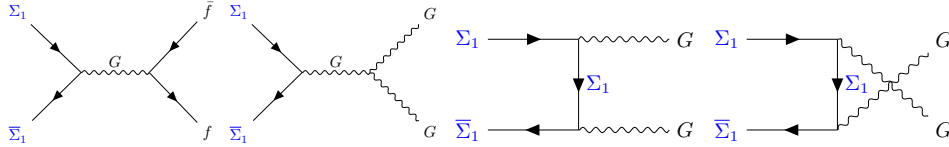
3. $\Delta L = 1$



4. $\Delta L = 2$



5. Gauge processes



Here f is the SM fermions and G is all possible Gauge Bosons.

Appendix D: Lepton asymmetry to baryon asymmetry conversion

1. Model-I

The asymmetry in the equilibrium number densities of particle n_i over antiparticle $\bar{n}_i$ can be written in the limit $\mu_i/T \ll 1$ as [90]

$$\begin{aligned} n_i - \bar{n}_i &= \frac{1}{6} g_i T^3 \left(\frac{\mu_i}{T} \right), \text{ fermion} \\ &= \frac{1}{3} g_i T^3 \left(\frac{\mu_i}{T} \right), \text{ boson} \end{aligned} \quad (\text{D1})$$

For a general derivation, we assume that the SM consists of N generations of quarks and leptons, m complex Higgs doublets, and we extended the SM by N' generations of Fermion triplets (Σ), one scalar triplet (Δ), one vector-like fermion (χ), and one singlet scalar (ϕ). The chemical potentials of the SM fields are assigned as follows: μ_W for W^- , μ_0 for $m \phi^0$ Higgs fields, μ_- for $m \phi^-$ Higgs fields, μ_{uL} for all left-handed up-quark fields,

μ_{dL} for all left-handed down-quark fields, μ_{uR} for all the right-handed up-quark fields, μ_{dR} for all the right-handed down-quark fields, μ_i for the left-handed neutrino fields, μ_{iL} for the left-handed charged lepton fields, and μ_{iR} for the right-handed charged lepton fields. The chemical potentials for the BSM fields are assigned as: μ_N for all N' N_R , μ_ρ for the ρ , μ_χ for χ , and μ_ϕ for ϕ . Now rapid interactions in the early universe enforces the following equilibrium relations among the chemical po-

tentials:

$$W^- \leftrightarrow \phi^- + \phi^0 \Rightarrow \mu_W = \mu_- + \mu_0 \quad (\text{D2})$$

$$W^- \leftrightarrow \bar{u}_L + d_L \Rightarrow \mu_{d_L} = \mu_W + \mu_{u_L} \quad (\text{D3})$$

$$W^- \leftrightarrow \bar{\nu}_{iL} + e_{iL} \Rightarrow \mu_{iL} = \mu_W + \mu_i \quad (\text{D4})$$

$$\phi^0 \leftrightarrow \bar{u}_L + u_R \Rightarrow \mu_{u_R} = \mu_0 + \mu_{u_L} \quad (\text{D5})$$

$$\phi^0 \leftrightarrow \bar{d}_R + d_L \Rightarrow \mu_{d_R} = -\mu_0 + \mu_W + \mu_{u_L} \quad (\text{D6})$$

$$\phi^0 \leftrightarrow e_{iL} + \bar{e}_{iR} \Rightarrow \mu_{iR} = -\mu_0 + \mu_W - \mu_i \quad (\text{D7})$$

$$\phi^{0*} + \phi^{0*} \leftrightarrow \rho + \rho^* \Rightarrow \mu_\rho = 0 \quad (\text{D8})$$

$$\begin{aligned} N_R \leftrightarrow \chi + \rho &\Rightarrow \mu_N = \mu_\rho + \mu_\chi \\ &\Rightarrow \mu_\chi = \mu_N \end{aligned} \quad (\text{D9})$$

$$\begin{aligned} \phi \leftrightarrow \bar{\chi} + \chi &\Rightarrow \mu_\phi = -\mu_\chi + \mu_\chi \\ &\Rightarrow \mu_\phi = 0 \end{aligned} \quad (\text{D10})$$

$$\nu_{iL} \leftrightarrow N_R + \phi^0 \Rightarrow \mu_i = \mu_N + \mu_0 \quad (\text{D11})$$

From Eq. (D11),

$$\begin{aligned} \sum_i \mu_i &= \sum_i \mu_N + \sum_i \mu_0 \\ &\Rightarrow \mu = N\mu_N + N\mu_0 \Rightarrow \mu = N\mu_\chi + N\mu_0 \\ &\Rightarrow \mu_\chi = \frac{\mu - N\mu_0}{N} \end{aligned} \quad (\text{D12})$$

Now the electroweak $B + L$ anomaly implies the existence of processes that correspond to the creation of $u_L d_L d_L \nu_L$ from each generation out of the vacuum. As long as these interactions are rapid, we have,

$$N(\mu_{u_L} + 2\mu_{d_L}) + \sum_i \mu_i = 0 \Rightarrow 3N\mu_{u_L} + 2N\mu_W + \mu = 0 \quad (\text{D13})$$

Let us now express the baryon (B), lepton (L), charge (Q), and third component of weak isospin (Q_3) number densities as,

$$\begin{aligned} B &= 3N\left(\frac{1}{3}\mu_{u_L} + \frac{1}{3}\mu_{u_R}\right) + 3N\left(\frac{1}{3}\mu_{d_L} + \frac{1}{3}\mu_{d_R}\right) \\ &= 4N\mu_{u_L} + 2N\mu_W \end{aligned} \quad (\text{D14})$$

$$\begin{aligned} L &= \sum_i (\mu_i + \mu_{iL} + \mu_{iR}) + \mu_\chi \\ &= 3\mu + 2N\mu_W - N\mu_0 + \mu_\chi \end{aligned} \quad (\text{D15})$$

$$\begin{aligned} Q &= 3N\left(\frac{2}{3}\mu_{u_L} - \frac{1}{3}\mu_{d_L} + \frac{2}{3}\mu_{u_R} - \frac{1}{3}\mu_{d_R}\right) \\ &\quad - \left(\sum_i \mu_{iL} + \sum_i \mu_{iR}\right) - 2 * 2\mu_W - 2m\mu_- \\ &= 2N\mu_{u_L} - 2\mu - (4N + 2m + 4)\mu_W + (4N + 2m)\mu_0 \end{aligned} \quad (\text{D16})$$

$$\begin{aligned} Q_3 &= 3N\left(\frac{1}{2}\mu_{u_L} - \frac{1}{2}\mu_{d_L}\right) + \sum_i \left(\frac{1}{2}\mu_i - \frac{1}{2}\mu_{iL}\right) - 2 * 2\mu_W \\ &\quad + 2m\left(\frac{1}{2}\mu_+ - \frac{1}{2}\mu_0\right) \\ &= -(2N + m + 4)\mu_W \end{aligned} \quad (\text{D17})$$

Now above the critical temperature both Q , and Q_3 are zero, which will give the B , and L in terms of μ_{u_L} as,

$$L = \mu_{u_L} N \left(\frac{4(N+1)}{m+2N} - 3 \left(\frac{1}{N} + 3 \right) \right) \quad (\text{D18})$$

$$B = 4N\mu_{u_L} \quad (\text{D19})$$

$$B = -\frac{4}{3\left(3 + \frac{1}{N}\right) - \frac{4(N+1)}{m+2N}} L \quad (\text{D20})$$

2. Model-II

In this case, the chemical potentials for the BSM fields are assigned as: μ_{Σ^0} for all $N' \Sigma^0$, μ_{Σ^+} for all Σ^+ , μ_δ for the δ^- , μ_{δ^0} for δ^0 , μ_χ for χ , and μ_ϕ for ϕ .

Now, rapid interactions in the early universe enforce the following equilibrium relations among the chemical potentials:

$$\phi^0 \leftrightarrow \bar{d}_R + d_L \Rightarrow \mu_{d_R} = -\mu_0 + \mu_W + \mu_{u_L} \quad (\text{D21})$$

$$\phi^0 \leftrightarrow e_{iL} + \bar{e}_{iR} \Rightarrow \mu_{iR} = -\mu_0 + \mu_W - \mu_i \quad (\text{D22})$$

$$\phi^0 \leftrightarrow \bar{u}_L + u_R \Rightarrow \mu_{u_R} = \mu_0 + \mu_{u_L} \quad (\text{D23})$$

$$W^- \leftrightarrow \phi^- + \phi^0 \Rightarrow \mu_W = \mu_- + \mu_0 \quad (\text{D24})$$

$$W^- \leftrightarrow \bar{u}_L + d_L \Rightarrow \mu_{d_L} = \mu_W + \mu_{u_L} \quad (\text{D25})$$

$$W^- \leftrightarrow \bar{\nu}_{iL} + e_{iL} \Rightarrow \mu_{iL} = \mu_W + \mu_i \quad (\text{D26})$$

$$\begin{aligned} W^- + W^- \leftrightarrow \delta^- + \delta^- &\Rightarrow \mu_W + \mu_W = \mu_\delta + \mu_\delta \\ &\Rightarrow \mu_\delta = \mu_W \end{aligned} \quad (\text{D27})$$

$$\begin{aligned} W^- + W^3 \leftrightarrow \delta^- + \delta^0 &\Rightarrow \mu_W + 0 = \mu_\delta + \mu_{\delta^0} \\ &\Rightarrow \mu_{\delta^0} = 0 \end{aligned} \quad (\text{D28})$$

$$\begin{aligned} \phi \leftrightarrow \bar{\chi} + \chi &\Rightarrow \mu_\phi = -\mu_\chi + \mu_\chi \\ &\Rightarrow \mu_\phi = 0 \end{aligned} \quad (\text{D29})$$

$$\Sigma^0 \leftrightarrow W^- + \Sigma^+ \Rightarrow \mu_{\Sigma^0} = \mu_W + \mu_{\Sigma^+} \quad (\text{D30})$$

$$\begin{aligned} \Sigma^0 \leftrightarrow \delta^0 + \chi &\Rightarrow \mu_{\Sigma^0} = \mu_{\delta^0} + \mu_\chi \\ &\Rightarrow \mu_\chi = \mu_{\Sigma^0} \end{aligned} \quad (\text{D31})$$

$$\nu_{iL} \leftrightarrow \Sigma^0 + \phi^0 \Rightarrow \mu_i = \mu_{\Sigma^0} + \mu_0 \quad (\text{D32})$$

From Eq. (D32),

$$\begin{aligned} \sum_i \mu_i &= \sum_i \mu_{\Sigma^0} + \sum_i \mu_0 \\ &\Rightarrow \mu = N\mu_{\Sigma^0} + N\mu_0 \Rightarrow \mu = N\mu_\chi + N\mu_0 \\ &\Rightarrow \mu_\chi = \frac{\mu - N\mu_0}{N} \end{aligned} \quad (\text{D33})$$

Now the electroweak $B + L$ anomaly implies the existence of processes that correspond to the creation of $u_L d_L d_L \nu_L$ from each generation out of the vacuum. As long as these interactions are rapid, we have,

$$N(\mu_{u_L} + 2\mu_{d_L}) + \sum_i \mu_i = 0 \Rightarrow 3N\mu_{u_L} + 2N\mu_W + \mu = 0 \quad (\text{D34})$$

Let us now express the baryon (B), lepton (L), charge (Q), and third component of weak isospin (Q_3) number densities as,

$$\begin{aligned} B &= 3N\left(\frac{1}{3}\mu_{u_L} + \frac{1}{3}\mu_{u_R}\right) + 3N\left(\frac{1}{3}\mu_{d_L} + \frac{1}{3}\mu_{d_R}\right) \\ &= 4N\mu_{u_L} + 2N\mu_W \end{aligned} \quad (\text{D35})$$

$$L = \sum_i (\mu_i + \mu_{iL} + \mu_{iR}) + \mu_\chi = 3\mu + 2N\mu_W - N\mu_0 + \mu_\chi \quad (\text{D36})$$

$$\begin{aligned} Q &= 3N\left(\frac{2}{3}\mu_{u_L} - \frac{1}{3}\mu_{d_L} + \frac{2}{3}\mu_{u_R} - \frac{1}{3}\mu_{d_R}\right) - \left(\sum_i \mu_{iL} + \sum_i \mu_{iR}\right) \\ &\quad - 2 * 2\mu_W - 2m\mu_- - 2\mu_\delta + N'\mu_{\Sigma^+} \\ &= 2N\mu_{u_L} - 2\mu - (4N + 2m + 6 - N')\mu_W \\ &\quad + (4N + 2m)\mu_0 + N'\mu_\chi \end{aligned}$$

$$\begin{aligned} Q_3 &= 3N\left(\frac{1}{2}\mu_{u_L} - \frac{1}{2}\mu_{d_L}\right) + \sum_i \left(\frac{1}{2}\mu_i - \frac{1}{2}\mu_{iL}\right) - 2 * 2\mu_W \\ &\quad + 2m\left(\frac{1}{2}\mu_+ - \frac{1}{2}\mu_0\right) - 2\mu_\delta + N'\mu_{\Sigma^+} \\ &= -(2N + m + 6 - N')\mu_W + N'\mu_\chi \end{aligned} \quad (\text{D38})$$

Now above the critical temperature both Q , and Q_3 are zero, which will give the B , and L in terms of μ_{u_L} as,

$$L = -\frac{2\mu_{u_L}(m^2(9N + 3) + m(32N^2 + 62N - 3N' + 18) + 2N(2N(7N + 22) - N' + 6))}{2m^2 + 8mN + 3m(N' + 4) + 2N(4N + N' + 12)} \quad (\text{D39})$$

$$B = \frac{8\mu_{u_L}N(m + 2N)(m + 2N + 6)}{2m^2 + 8mN + 3m(N' + 4) + 2N(4N + N' + 12)} \quad (\text{D40})$$

$$B = -\frac{4N(m + 2N)(m + 2N + 6)}{m^2(9N + 3) + m(32N^2 + 62N - 3N' + 18) + 2N(2N(7N + 22) - N' + 6)}L \quad (\text{D41})$$

Appendix E: Dark matter decay width

$$\begin{aligned} \Gamma_\chi^{(1)}(\chi \rightarrow eW) &= \frac{G_\chi^{(1)2} M_\chi^3}{16\pi M_W^2} \left(1 + \frac{m_e^4}{M_\chi^4} - \frac{2M_W^4}{M_\chi^4} - \frac{2m_e^2}{M_\chi^2} + \frac{M_W^2}{M_\chi^2} + \frac{m_e^2 M_W^2}{M_\chi^4} \right) \\ &\quad \times \left(1 + \frac{m_e^4}{M_\chi^4} + \frac{M_W^4}{M_\chi^4} - \frac{2m_e^2}{M_\chi^2} - \frac{2M_W^2}{M_\chi^2} - \frac{2m_e^2 M_W^2}{M_\chi^4} \right)^{\frac{1}{2}}, \end{aligned} \quad (\text{E1})$$

$$\begin{aligned} \Gamma_\chi^{(2)}(\chi \rightarrow \nu Z) &= \frac{G_\chi^{(2)2} M_\chi^3}{16\pi M_Z^2} \left(1 + \frac{m_\nu^4}{M_\chi^4} - \frac{2M_Z^4}{M_\chi^4} - \frac{2m_\nu^2}{M_\chi^2} + \frac{M_Z^2}{M_\chi^2} + \frac{m_\nu^2 M_Z^2}{M_\chi^4} \right) \\ &\quad \times \left(1 + \frac{m_\nu^4}{M_\chi^4} + \frac{M_Z^4}{M_\chi^4} - \frac{2m_\nu^2}{M_\chi^2} - \frac{2M_Z^2}{M_\chi^2} - \frac{2m_\nu^2 M_Z^2}{M_\chi^4} \right)^{\frac{1}{2}}, \end{aligned} \quad (\text{E2})$$

$$\Gamma_\chi^{(3)}(\chi \rightarrow \phi\nu) = \frac{G_\chi^{(3)2} M_\chi}{16\pi} \left(1 + \frac{m_\nu^2}{M_\chi^2} - \frac{M_\phi^2}{M_\chi^2} + \frac{2m_\nu}{M_\chi} \right) \left(1 + \frac{m_\nu^4}{M_\chi^4} + \frac{M_\phi^4}{M_\chi^4} - \frac{2m_\nu^2}{M_\chi^2} - \frac{2M_\phi^2}{M_\chi^2} - \frac{2m_\nu^2 M_\phi^2}{M_\chi^4} \right)^{\frac{1}{2}}, \quad (\text{E3})$$

$$\Gamma_{\chi}^{(4)}(\chi \rightarrow \phi \nu) = \frac{G_{\chi}^{(4)2} M_{\chi}}{16\pi} \left(1 + \frac{m_{\nu}^2}{M_{\chi}^2} - \frac{M_{\phi}^2}{M_{\chi}^2} + \frac{2m_{\nu}}{M_{\chi}} \right) \left(1 + \frac{m_{\nu}^4}{M_{\chi}^4} + \frac{M_{\phi}^4}{M_{\chi}^4} - \frac{2m_{\nu}^2}{M_{\chi}^2} - \frac{2M_{\phi}^2}{M_{\chi}^2} - \frac{2m_{\nu}^2 M_{\phi}^2}{M_{\chi}^4} \right)^{\frac{1}{2}}, \quad (\text{E4})$$

$$\begin{aligned} \Gamma_{\chi}^{(5)}(\chi \rightarrow \nu f \bar{f}) = & \frac{G_{\chi}^{(5)2} M_{\chi}^5}{192\pi^3} \left(\frac{1}{4} (1 - 4s_w^2 + 8s_w^4) \left((1 - 14 \frac{m_f^2}{M_{\chi}^2} - 2 \frac{m_f^4}{M_{\chi}^4} - 12 \frac{m_f^6}{M_{\chi}^6}) \sqrt{1 - 4 \frac{m_f^2}{M_{\chi}^2}} \right. \right. \\ & + 12 \frac{m_f^4}{M_{\chi}^4} \left(\frac{m_f^4}{M_{\chi}^4} - 1 \right) \log \left[\frac{1 - 3 \frac{m_f^2}{M_{\chi}^2} - (1 - \frac{m_f^2}{M_{\chi}^2}) \sqrt{1 - 4 \frac{m_f^2}{M_{\chi}^2}}}{\frac{m_f^2}{M_{\chi}^2} (1 + \sqrt{1 - 4 \frac{m_f^2}{M_{\chi}^2}})} \right] \Big) \\ & + 2s_w^2 (2s_w^2 - 1) \left(\frac{m_f^2}{M_{\chi}^2} (2 + 10 \frac{m_f^2}{M_{\chi}^2} - 12 \frac{m_f^4}{M_{\chi}^4}) \sqrt{1 - 4 \frac{m_f^2}{M_{\chi}^2}} \right. \\ & \left. \left. + 6 \frac{m_f^4}{M_{\chi}^4} (1 - 2 \frac{m_f^2}{M_{\chi}^2} + 2 \frac{m_f^4}{M_{\chi}^4}) \log \left[\frac{1 - 3 \frac{m_f^2}{M_{\chi}^2} - (1 - \frac{m_f^2}{M_{\chi}^2}) \sqrt{1 - 4 \frac{m_f^2}{M_{\chi}^2}}}{\frac{m_f^2}{M_{\chi}^2} (1 + \sqrt{1 - 4 \frac{m_f^2}{M_{\chi}^2}})} \right] \right) \right), \quad (\text{E5}) \end{aligned}$$

where the effective couplings $G_{\chi}^{(1)}$, $G_{\chi}^{(2)}$, $G_{\chi}^{(3)}$, $G_{\chi}^{(4)}$, and $G_{\chi}^{(5)}$ can be given as

$$G_{\chi}^{(1)} = (y_{\chi} v_1) \left(\frac{1}{M_{\Sigma_1}} \right) \frac{(y_{\Sigma} v_0)}{\sqrt{2}} \left(\frac{1}{M_{\chi}} \right) \left(\frac{g}{2 \cos \theta_w} \right), \quad (\text{E6})$$

$$G_{\chi}^{(2)} = (y_{\chi} v_1) \left(\frac{1}{M_{\Sigma_1}} \right) \frac{(y_{\Sigma} v_0)}{\sqrt{2}} \left(\frac{1}{M_{\chi}} \right) \left(\frac{g}{\sqrt{2}} \right), \quad (\text{E7})$$

$$G_{\chi}^{(3)} = \lambda_{\text{DM}} \left(\frac{1}{M_{\chi}} \right) (y_{\chi} v_1) \left(\frac{1}{M_{\Sigma_1}} \right) \frac{(y_{\Sigma} v_0)}{\sqrt{2}}, \quad (\text{E8})$$

$$G_{\chi}^{(4)} = (y_{\chi} v_1) \left(\frac{1}{M_{\Sigma_1}} \right) (y_{\Sigma}) (\sin(\gamma)), \quad (\text{E9})$$

$$G_{\chi}^{(5)} = (y_{\chi} v_1) \left(\frac{1}{M_{\Sigma_1}} \right) (y_{\Sigma}) \left(\frac{1}{M_{h_1}^2} \right) \left(\frac{m_f}{v_0} \right), \quad (\text{E10})$$

where v_1 is the VEV of ρ (Δ).

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